Colloque de prospective du PNPS 2018

The key role of the host star for the evolution of planetary systems

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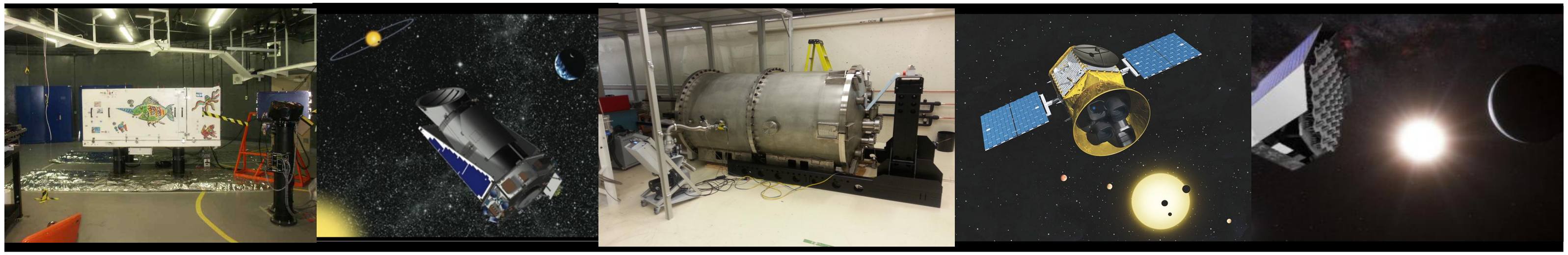
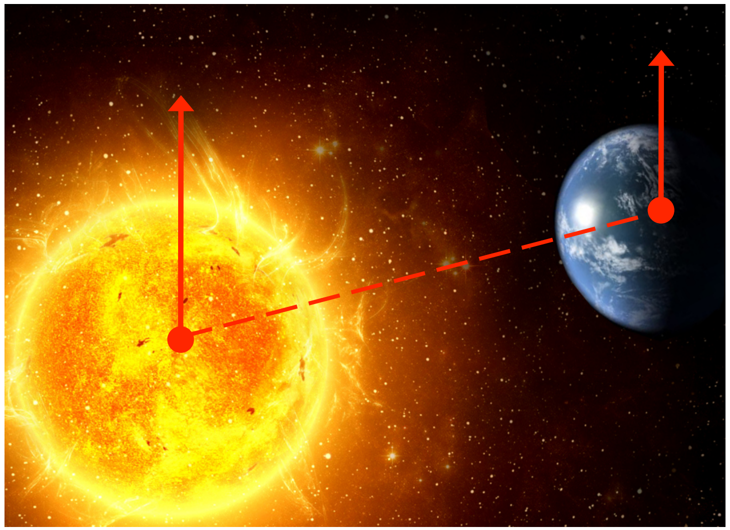
The general context

A revolution in astrophysics: discovery of **new planetary systems**
& characterisation of **the dynamics of their host (multiple) stars**
(asteroseismology and spectropolarimetry)



ESPaDOnS@CFHT LPs — *Kepler – K2* — *SPIRou* — *CHEOPS & TESS* — *PLATO*

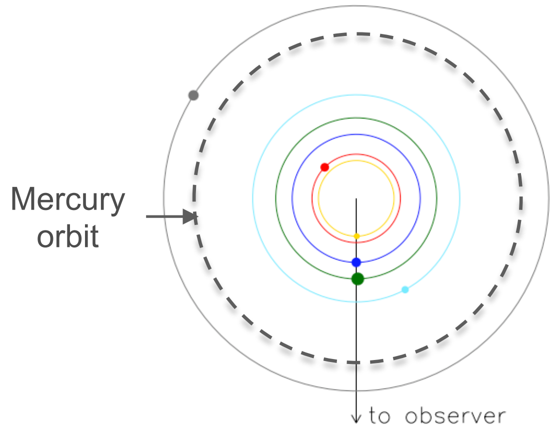
Stellar rotation & magnetism/activity
– planetary dynamics/atmospheres



Albrecht et al. 2012; Gizon,…, Mathis,…, et al. 2013

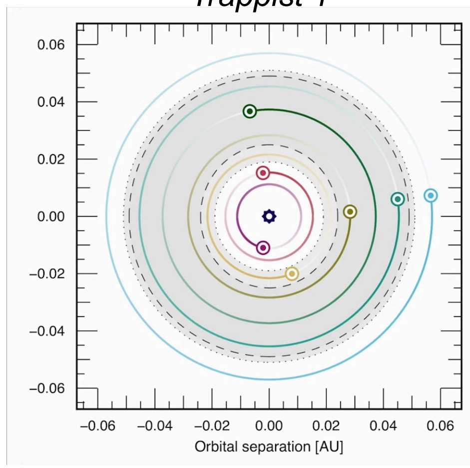
Orbital architecture
→ Interacting systems (tides/strong winds/radiation)

Kepler 11



Lissauer et al. 2011
Bolmont et al. 2014

Trappist 1



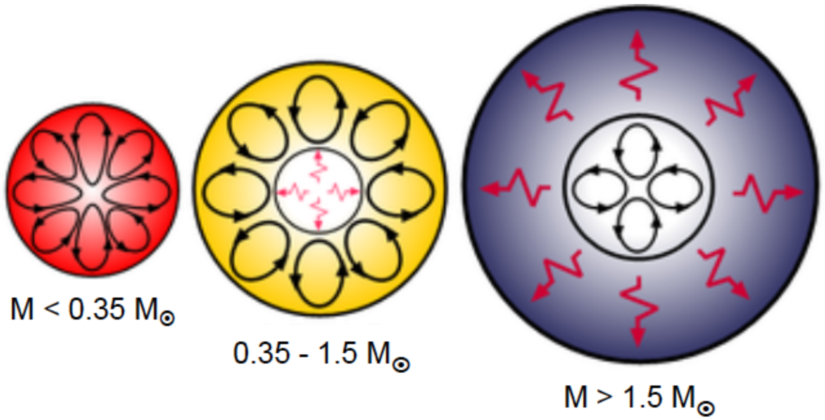
Guillon et al. 2017

State of the art

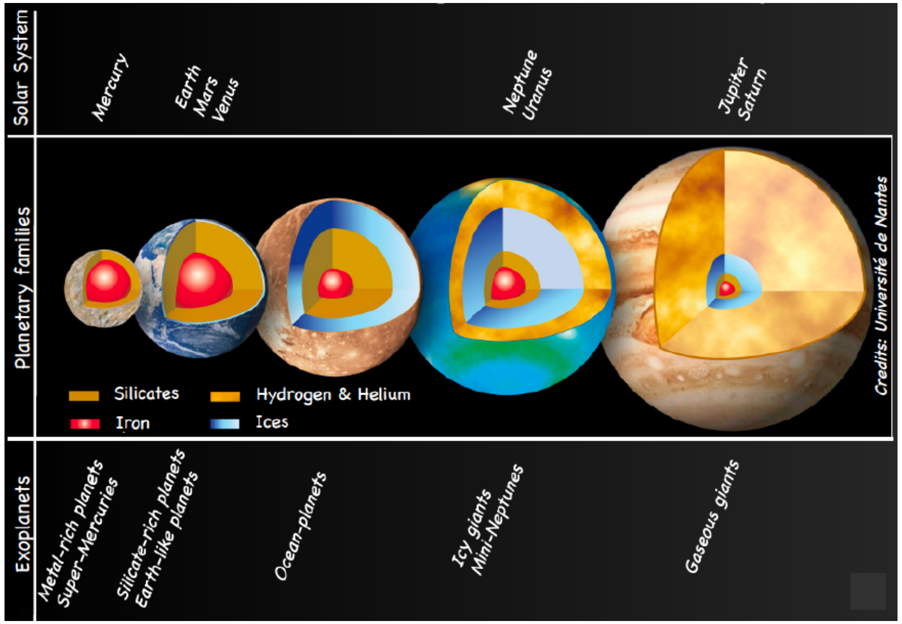
In studies of star-star or -planet systems, bodies are treated as point-mass objects or solids with ad-hoc prescriptions for tides, stellar winds and electromagnetic interactions

However their **complex internal structure, evolution, rotation, and magnetism** impact **tidal (and magnetic) Star-Planet Interactions**

Host star (M in $M_\odot$)



Planets



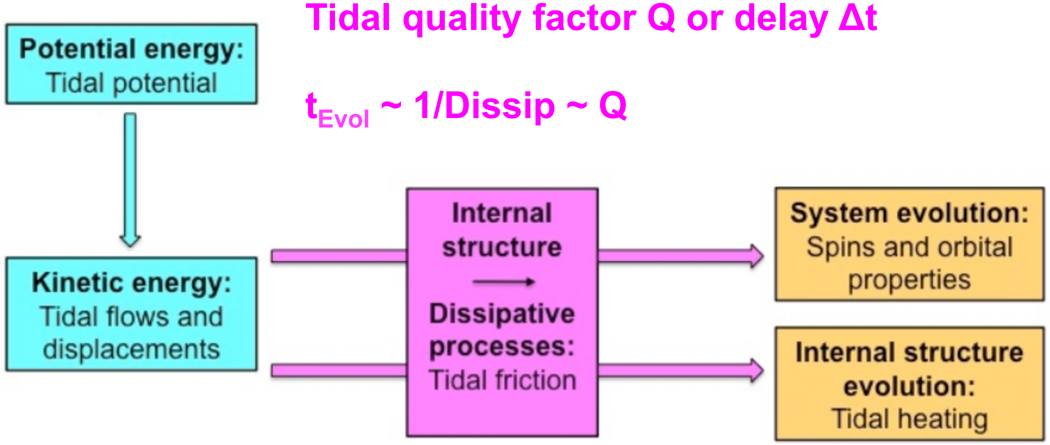
→ Need of an ab-initio physical modeling to accompany the study of discovered systems

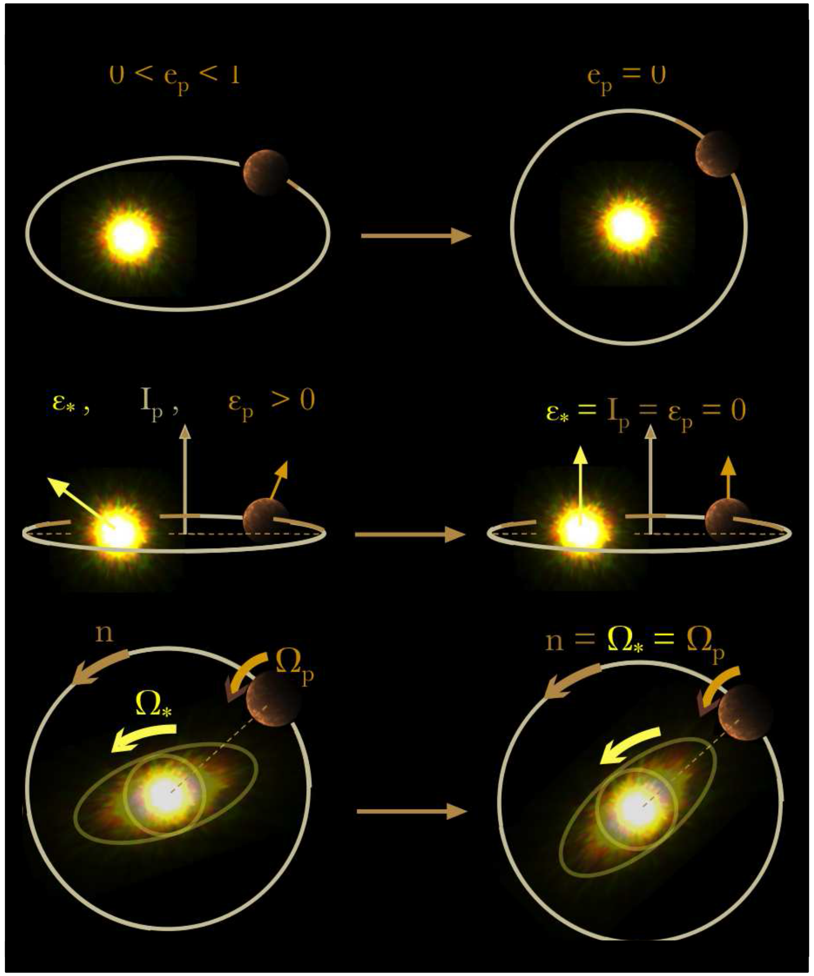
The “engine” of the tidal evolution of binary systems: friction & energy dissipation

Dynamical evolution of a binary system

Tidal quality factor Q or delay Δt

t_{Evol} ~ 1/Dissip ~ Q





Mathis & Remus 2013

→ Necessity to identify the dissipative processes and to evaluate their strength along the evolution of systems and of their components

→ Time-scales for circularization, synchronization, alignment, and migration (→ Age)

Tidal velocities/displacements

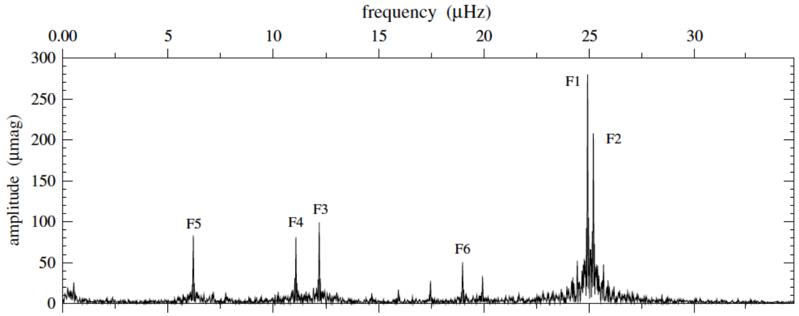
In stars and fluid planetary layers:

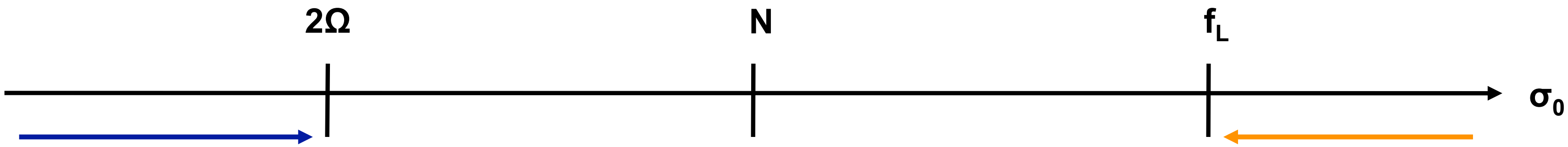
- Large-scale circulation resulting from the hydrostatic adjustment to the tidal perturbation (*equilibrium tide*)

- Waves excited by the tidal potential (*dynamical tide*)

Excitation by each Fourier component of the forcing

KOI-54: Welsh et al. 2011





Inertial waves (C.Z.)

Acoustic waves

Gravito-inertial waves (Stab. strat.)

Mathis & Remus 2013

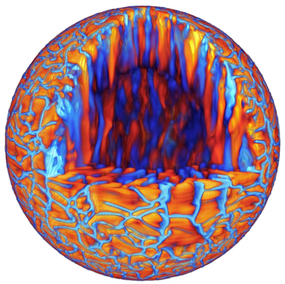
+ Alfvén waves if magnetic fields

Dissipative mechanisms:

- **Convective regions:** turbulence

- **Stably stratified regions**: heat diffusion

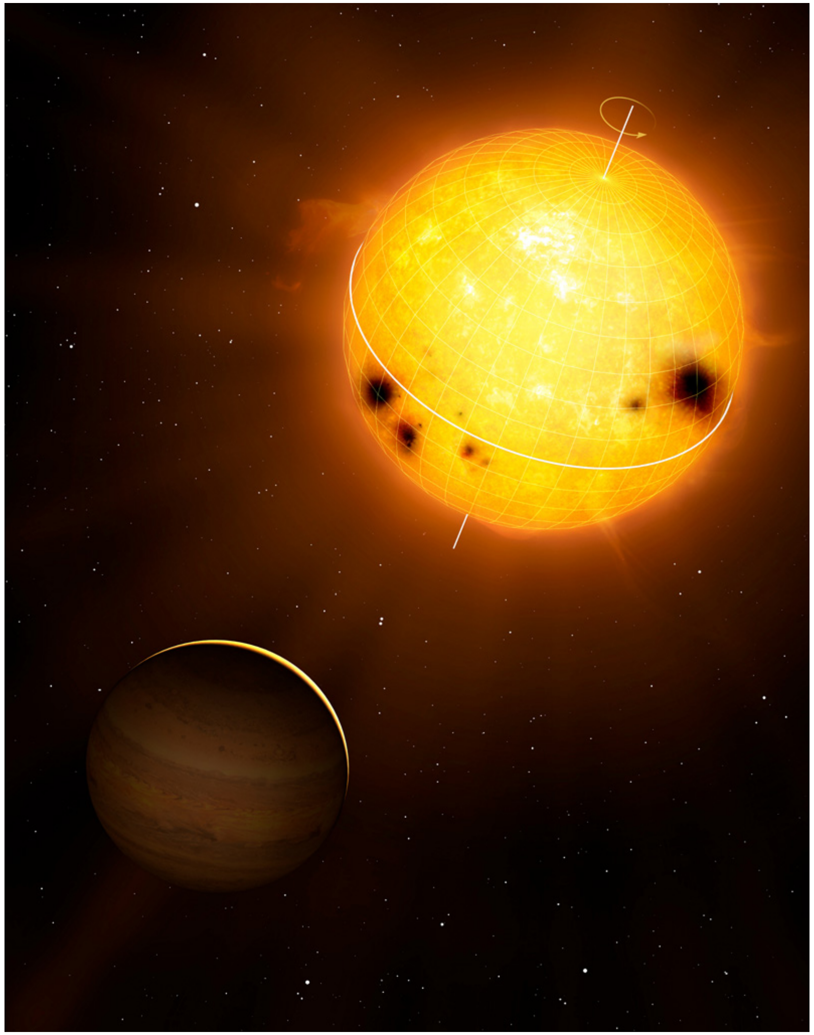
Brun et al.



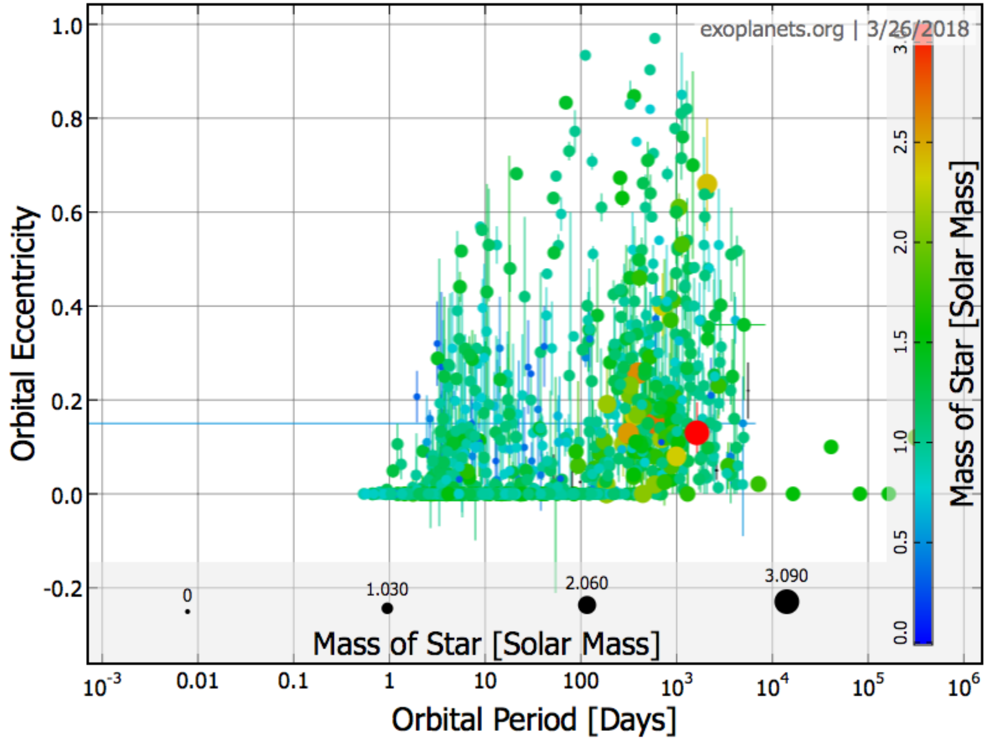
+ Ohmic heating if magnetic fields

The signature of tidal interactions in exoplanetary systems & multiple stars

The case of hot-Jupiter systems (and binary solar-type stars)



Gizon et al. 2013; Davies et al. 2015



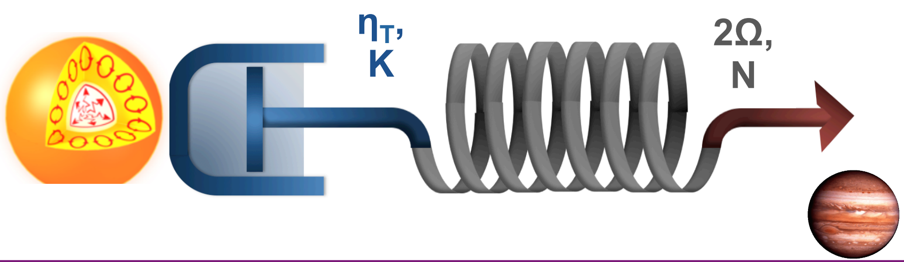
→ Tidal dissipation in a star varies over **several orders of magnitude** as a function of:

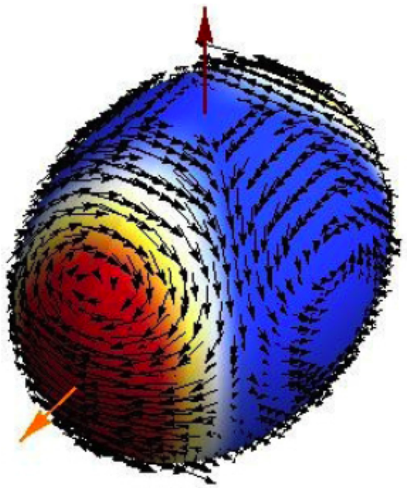
- The mass
- The age
- The dynamics (rotation)

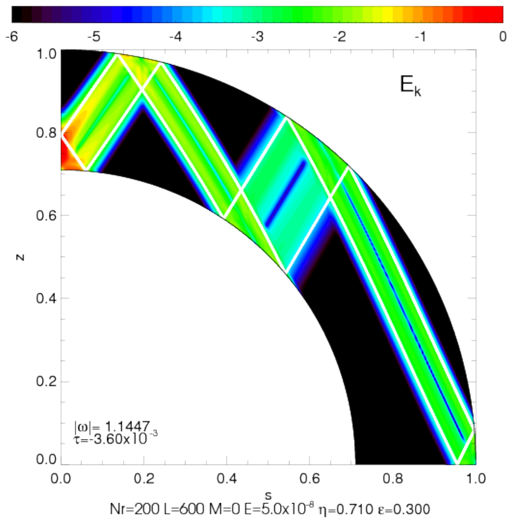
→ need for ab-initio modeling

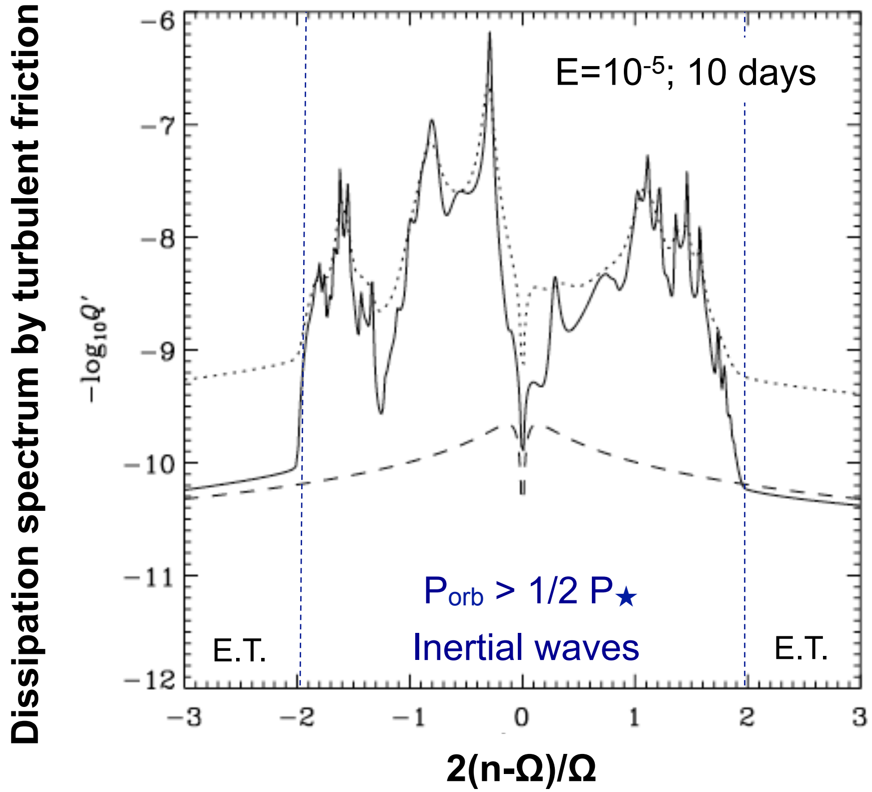
Tidal dissipation in low-mass star convective envelopes

Fluid dynamical tide: forced (gravito-) inertial waves → resonant response







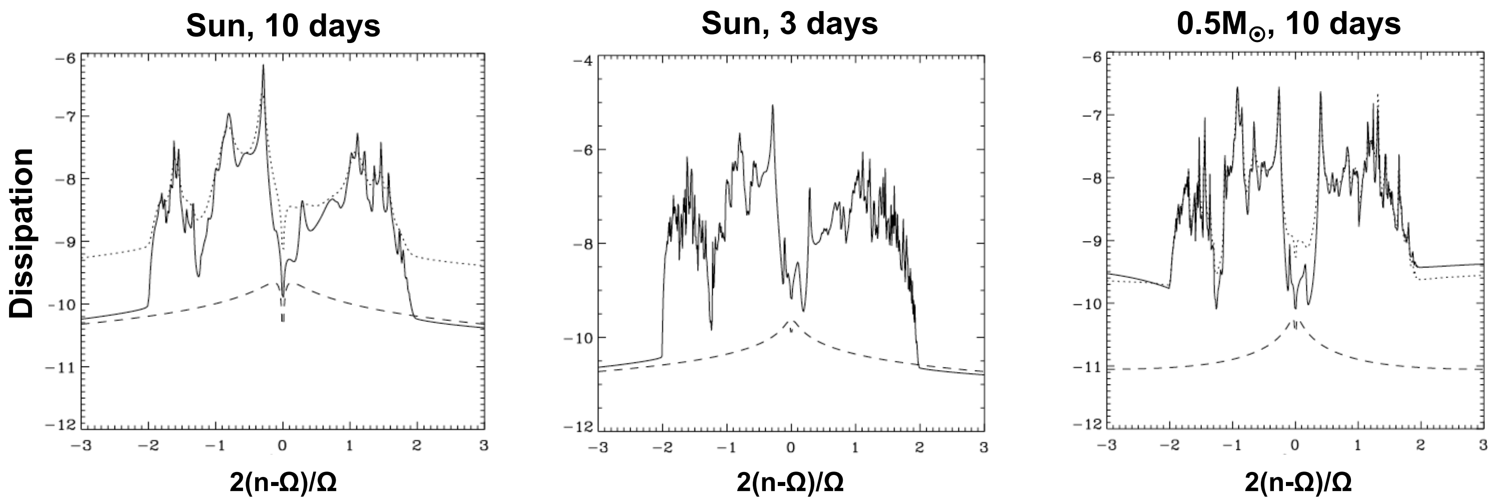


Ogilvie & Lin 2004, 2007
Rieutord & Valdetarro 2010
Baruteau & Rieutord 2013
Guenel et al. 2016
Lin & Ogilvie 2017
Wei 2018

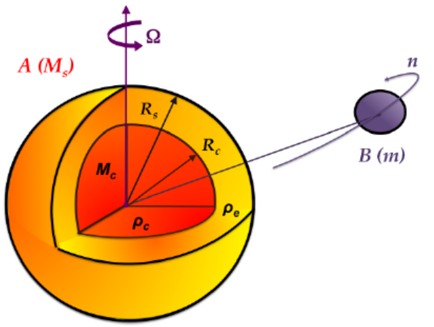
Dissipation variations with stellar parameters

As a function of stellar mass, age and rotation

Ogilvie & Lin 2007



To get an order of magnitude of tidal dissipation along the evolution of stars
→ a frequency-averaged dissipation (and the equivalent tidal quality factor)



$$\frac{3}{2\overline{Q'}} = \frac{\overline{k_2}}{Q} = \int_{-\infty}^{+\infty} \mathrm{Im}\left[k_2^2(\omega)\right]\frac{d\omega}{\omega} = \left\langle \mathrm{Im}\left[k_2^2(\omega)\right]\right\rangle_\omega = \frac{100\pi}{63}\epsilon^2\left(\frac{\alpha^5}{1-\alpha^5}\right)(1-\gamma)^2$$

$$\times(1-\alpha)^4\left(1+2\alpha+3\alpha^2+\frac{3}{2}\alpha^3\right)^2\left[1+\left(\frac{1-\gamma}{\gamma}\right)\alpha^3\right]\left[1+\frac{3}{2}\gamma+\frac{5}{2\gamma}\left(1+\frac{1}{2}\gamma-\frac{3}{2}\gamma^2\right)\alpha^3-\frac{9}{4}(1-\gamma)\alpha^5\right]^{-2}$$

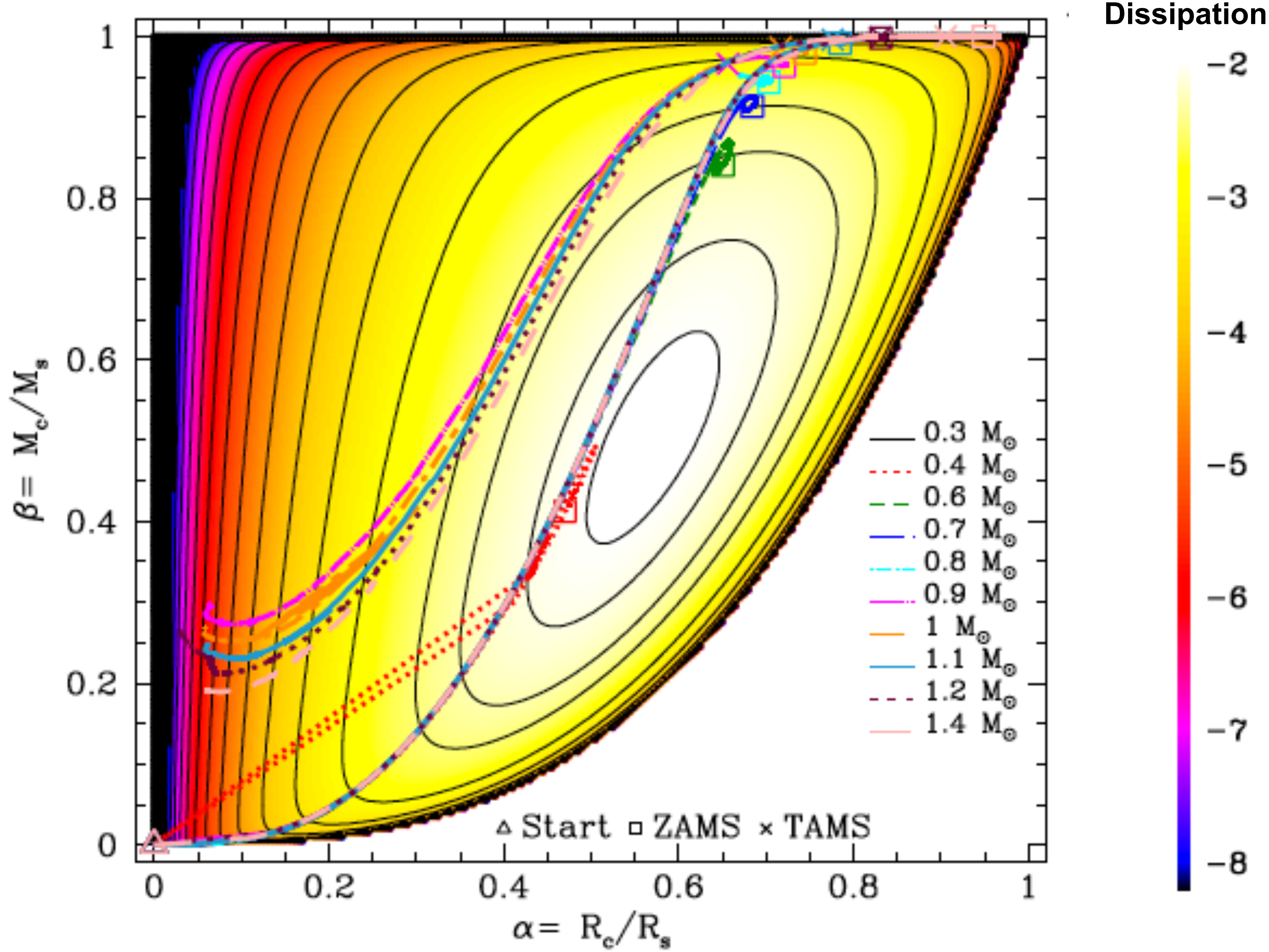
with

$$\alpha = \frac{R_c}{R_s},\quad \beta = \frac{M_c}{M_s}\quad \text{and}\quad \gamma = \frac{\rho_e}{\rho_c} = \frac{\alpha^3(1-\beta)}{\beta(1-\alpha^3)} < 1. \qquad \text{structure}$$

$$\epsilon^2 \equiv \left(\Omega/\sqrt{\mathcal{G}M_s/R_s^3}\right)^2 = (\Omega/\Omega_c)^2 \ll 1 \qquad \text{rotation}$$

Ogilvie 2013; Mathis 2015

The tidal H-R diagram



Mathis 2015; Gallet et al. 2017; Bolmont et al. 2017

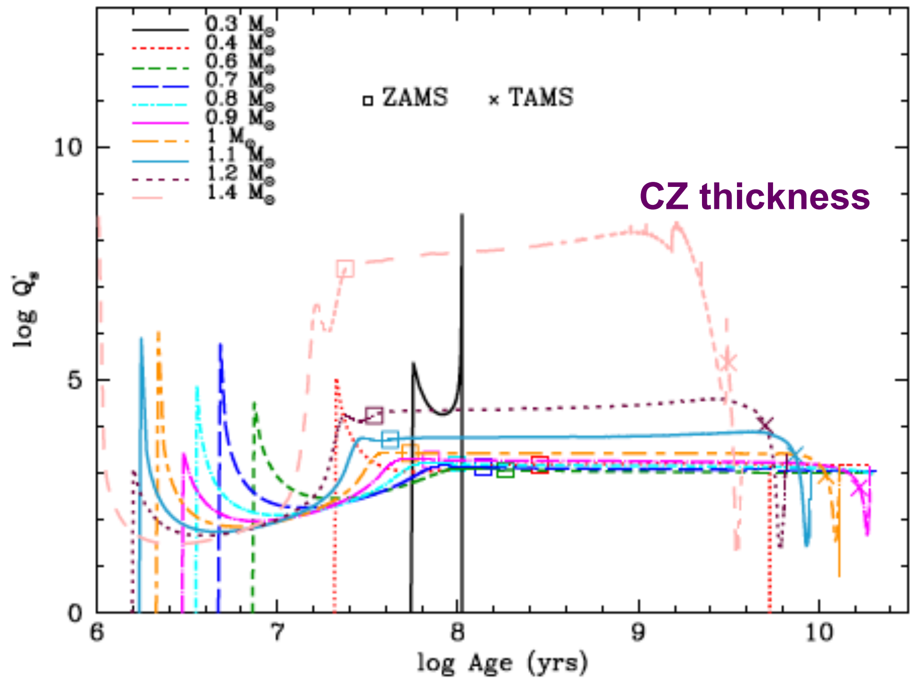
Grids of tidal dissipation for star-planet and multiple star systems

In low-mass and solar-type stars, it varies over several orders of magnitude:

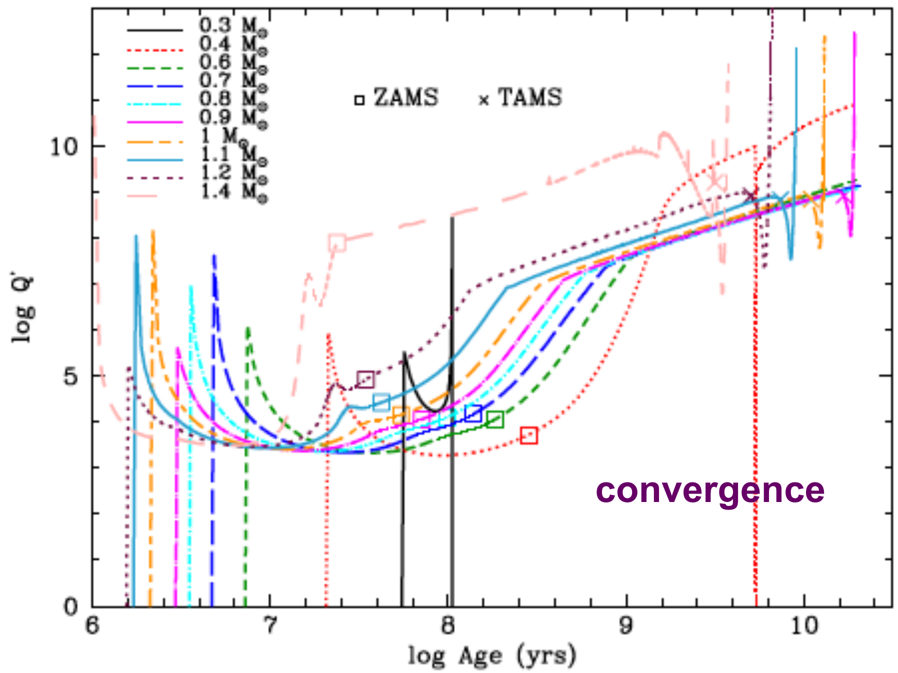
→ Stronger dynamical tide along the Pre-Main-Sequence and Sub-Giant phases

→ Its amplitude is driven by the structural evolution on the PMS and the rotational evolution on the MS

→ Necessity to **couple structural and rotational evolutions**



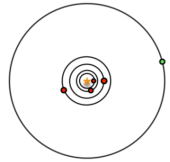
Structural evolution



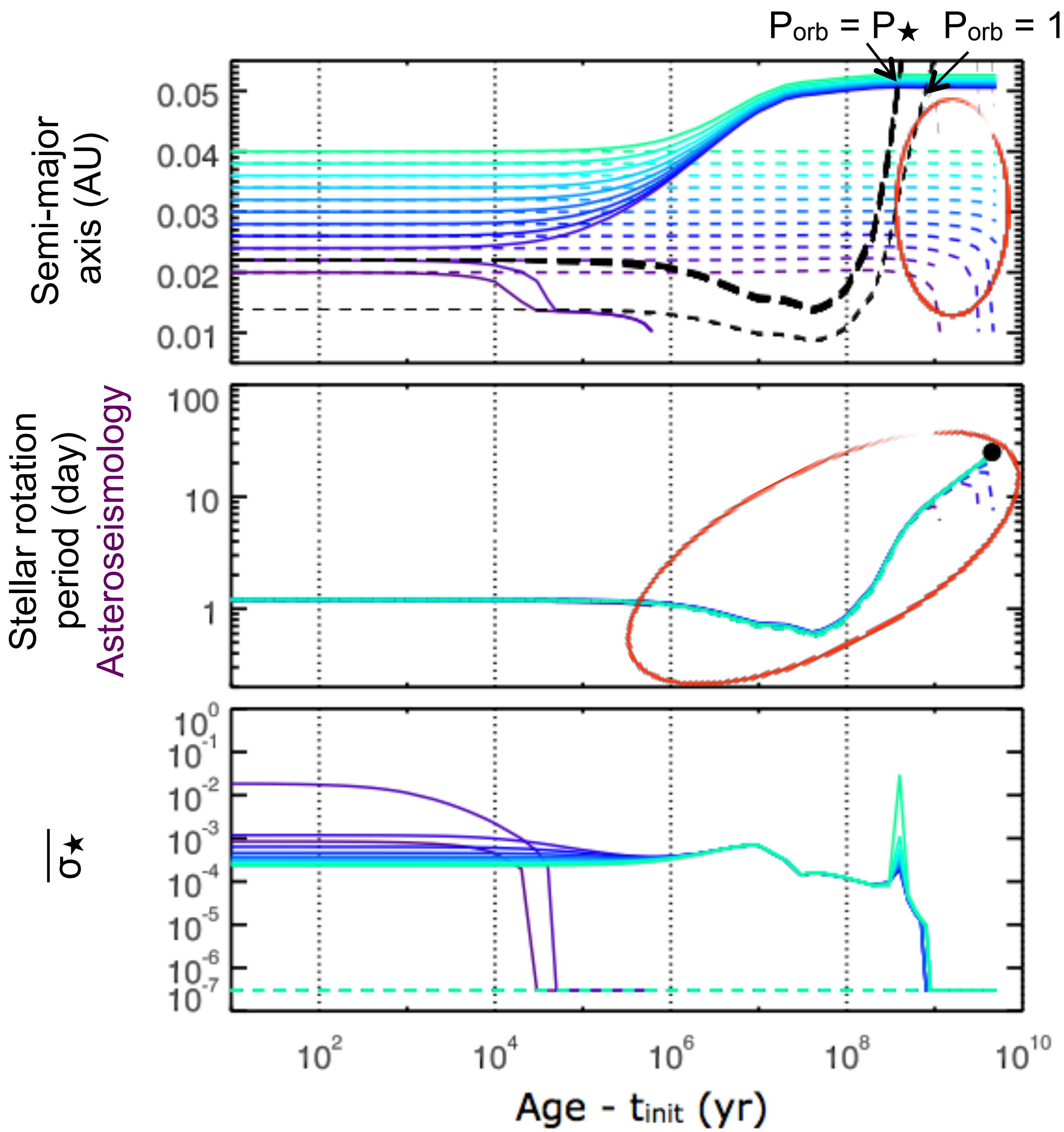
Structural & rotational evolution

Mathis 2015; Gallet et al. 2017; Bolmont et al. 2017

Star-planet systems tidal evolution



- Low-mass star-planet systems - circular & coplanar
- Ab-initio frequency-averaged dissipation of stellar tides in the convective envelope



Standard model - - -

Model Bolmont & Mathis ——

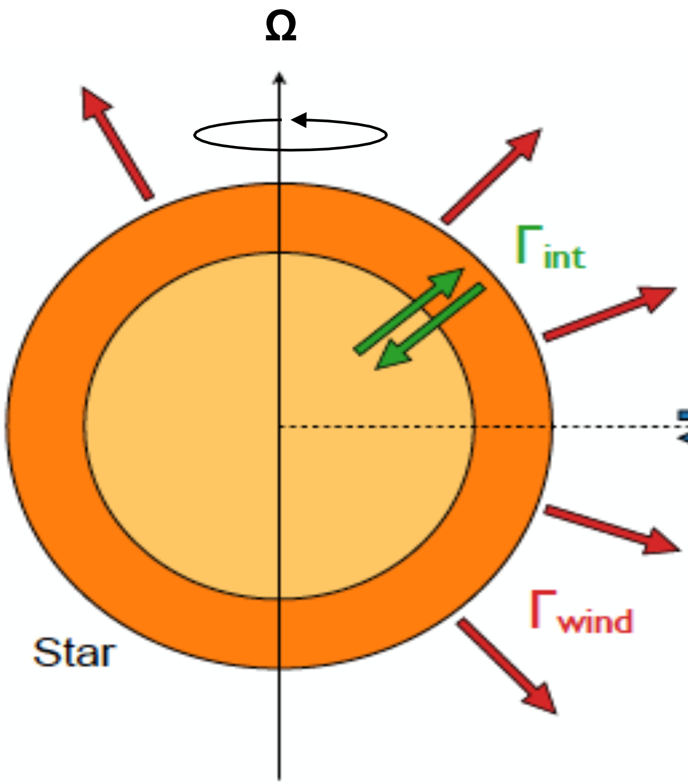
$M_p = 1\ M_{jup}$

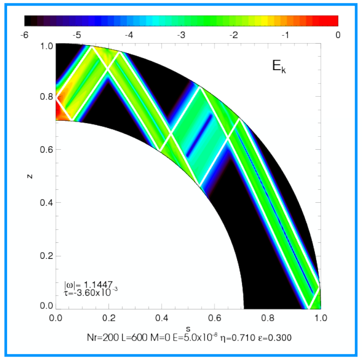
$M_\star = 1\ M_\odot$

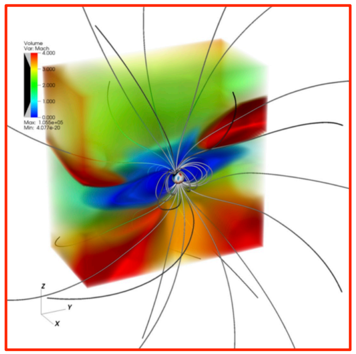
$P_{\star,i} = 1.2$ day

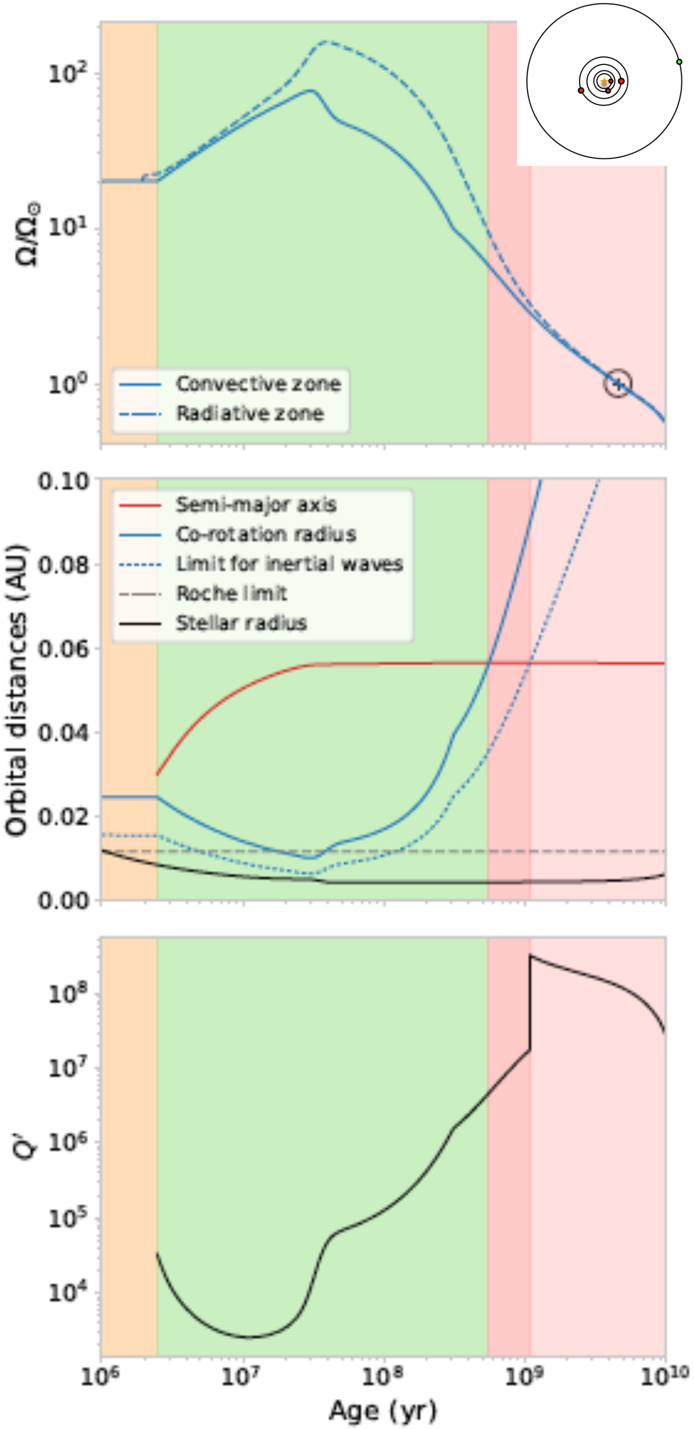
Bolmont & Mathis 2016

Impact on stellar rotation (I)







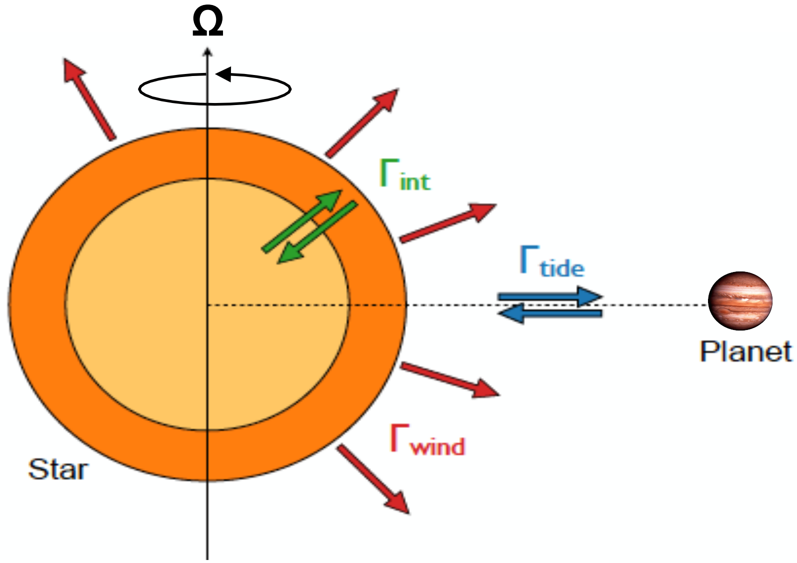


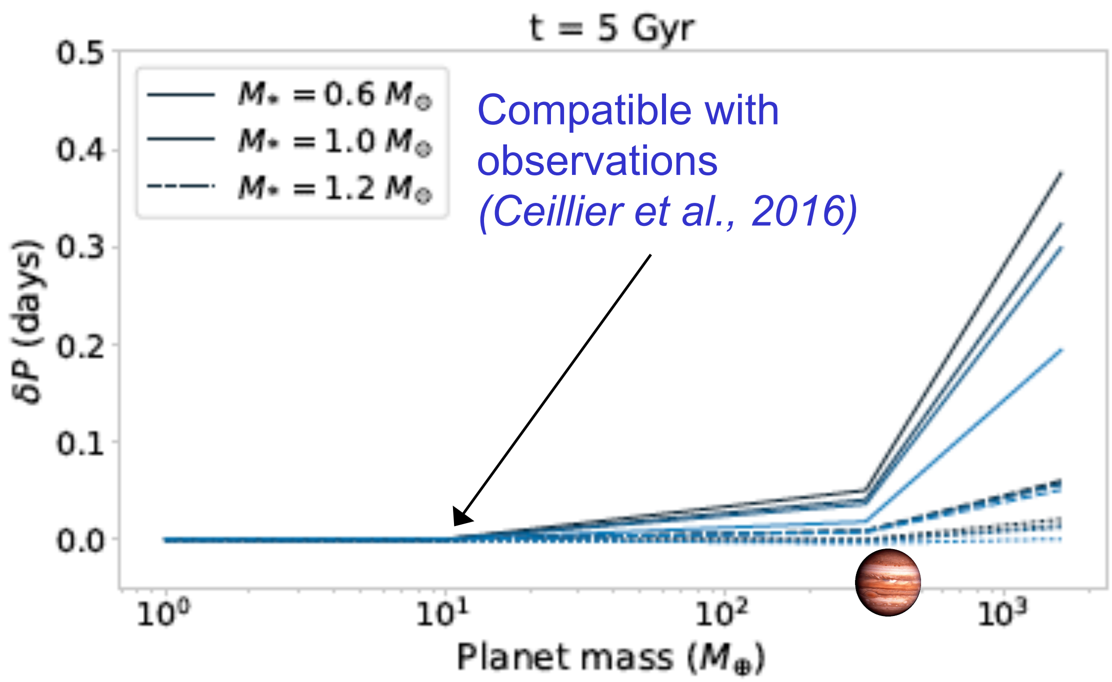
Benbakoura et al. 2018

Mathis 2015 *Reville et al. 2016*

ESPEM: Evolution Systèmes Planétaires Et Magnétisme

Impact on stellar rotation (II)





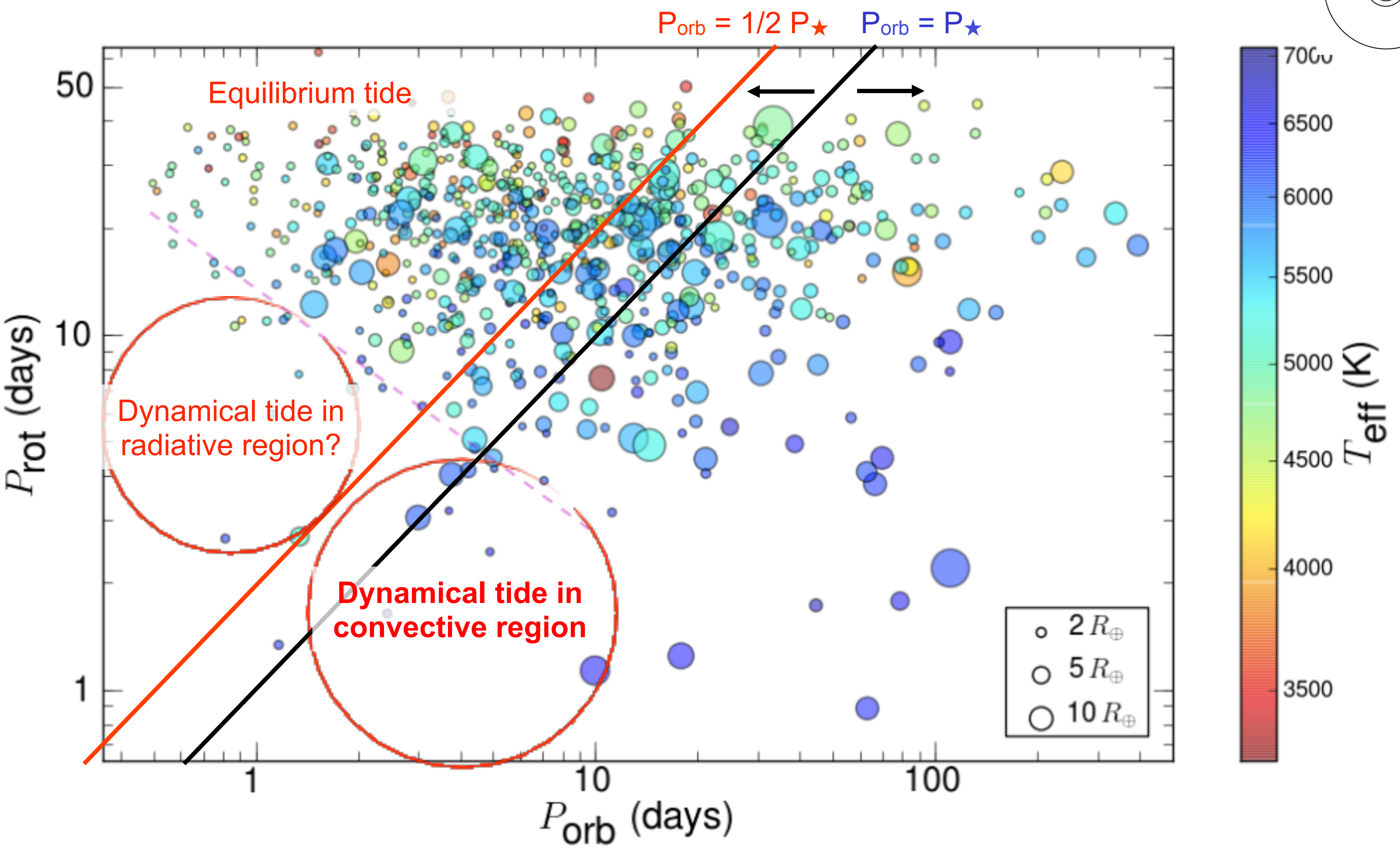
Planets migrate inward/outward
→ the star spins up/down

➡ Rotation excess: star initially slow rotator

➡ Rotation deficiency: star initially fast rotator

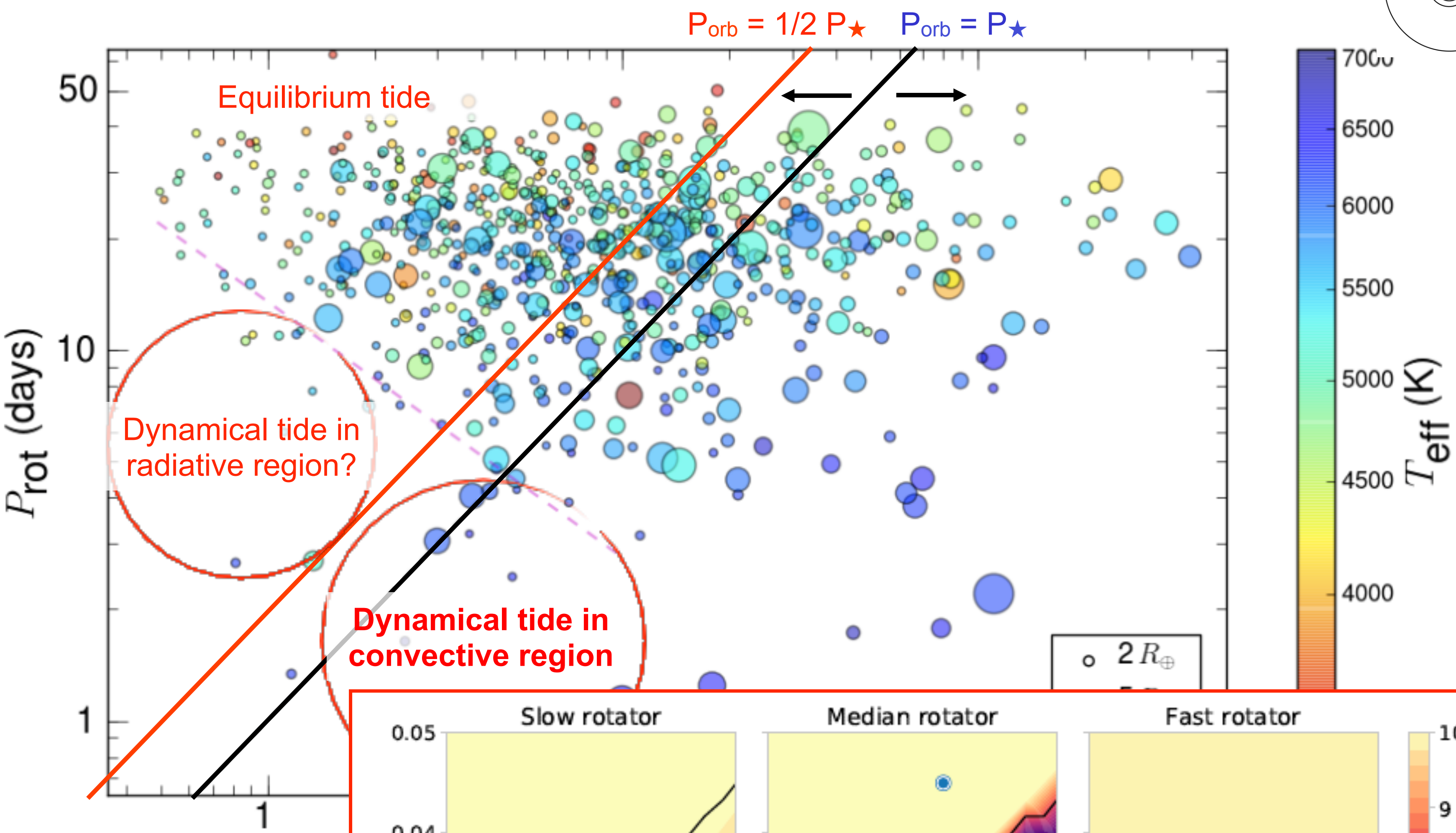
Bolmont & Mathis 2016; Benbakoura et al. 2018

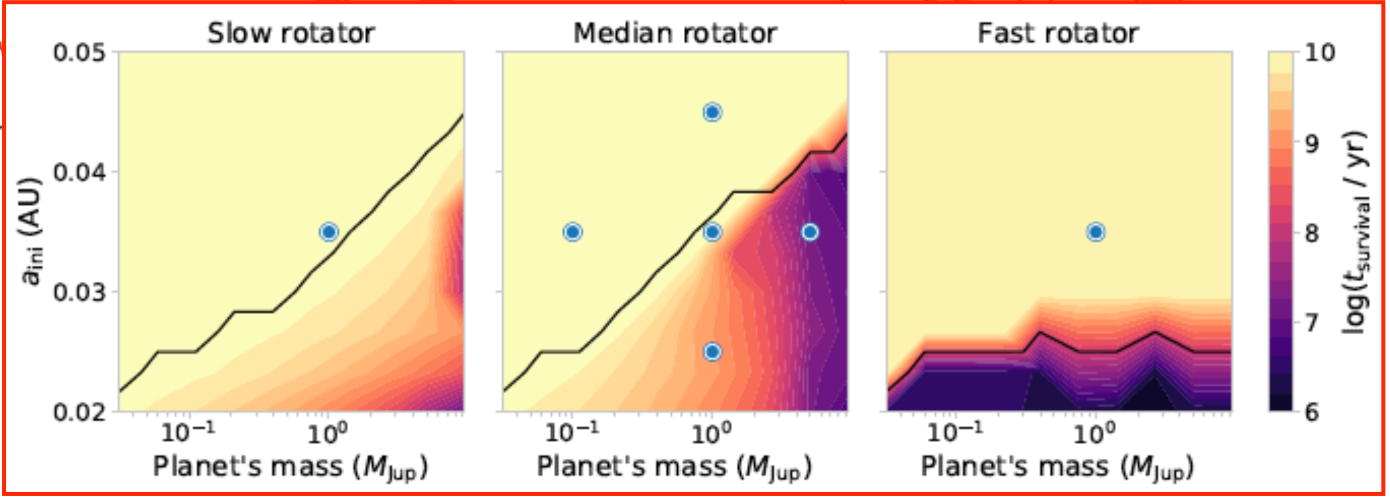
Understanding hot-Jupiters systems



McQuillan et al. 2013

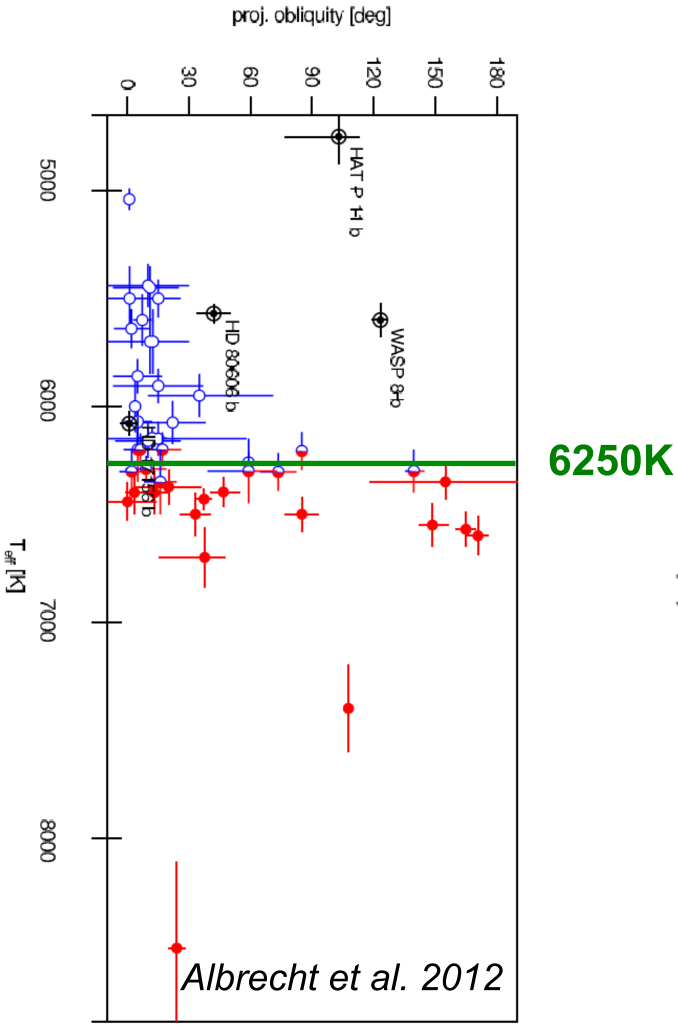
Understanding hot-Jupiters systems

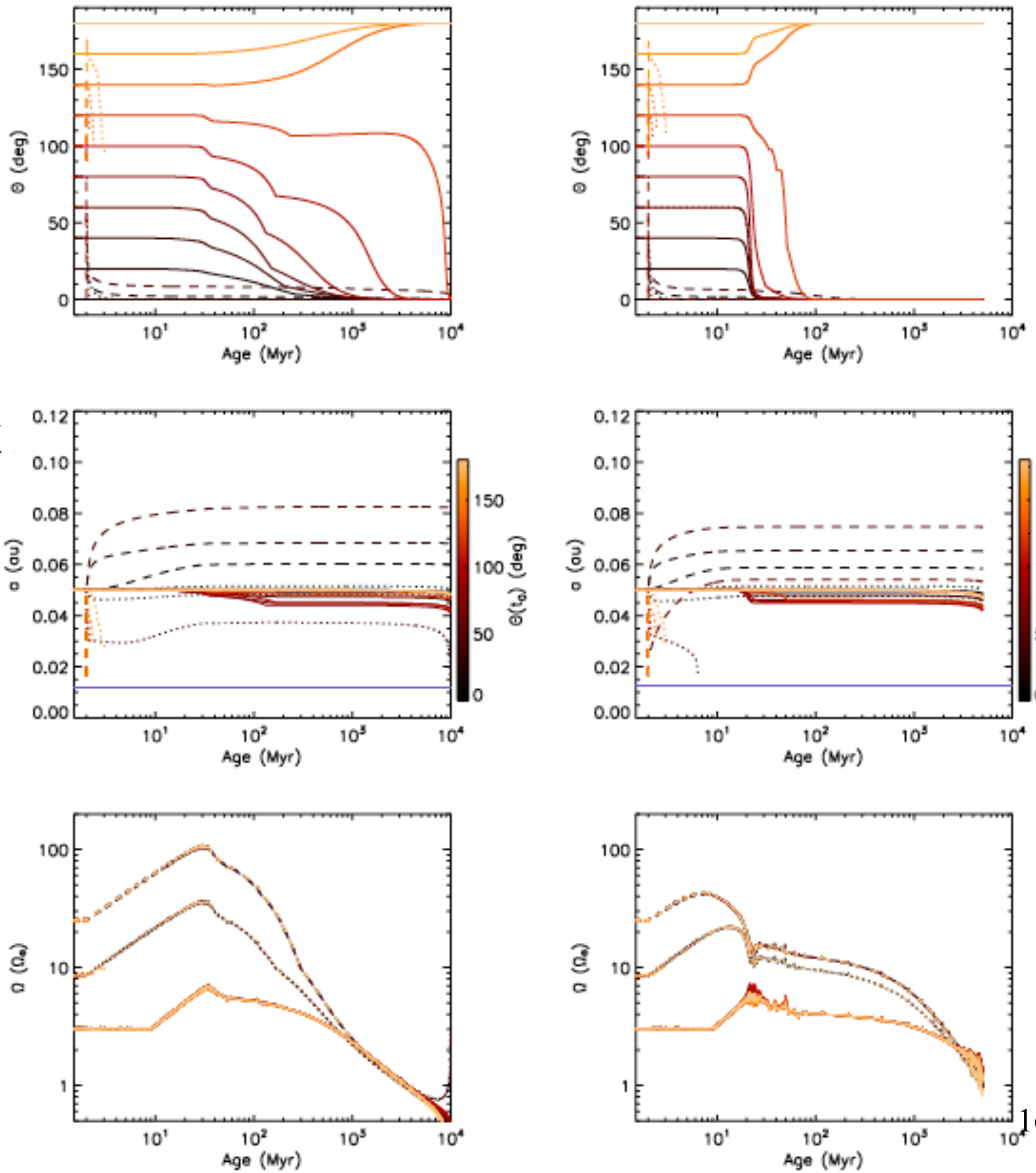




Benbakoura et al. 2018

Understanding the spin-orbit angles





Tides are too efficient to damp spin-orbit angles

Damiani & Mathis 2018

Conclusions and perspectives

Summary:

- Tidal dissipation in stellar convective zones varies over several orders of magnitude as a function of stellar mass, age and rotation
- The dynamical tide causes a much faster evolution than the equilibrium tide

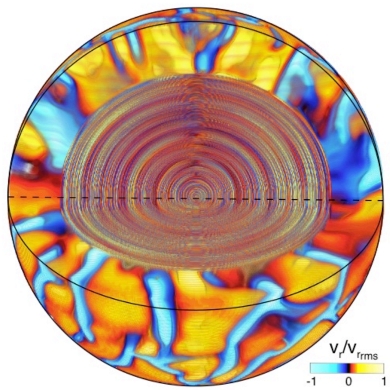
  → Needs to be taken into account in tidal studies
  → Implications on the understanding of planets distribution

- The dynamical tide is strong enough so that the star's early rotation history has a strong influence on close-in planets
- For $M_p > 10\ M_\oplus$, the dynamical tide induced migration is strong enough to influence the star's rotation

Perspectives:

Treat:
→ Multiple systems
→ Eccentric orbits and inclined systems

Take into account:
→ Tidal dissipation frequency-dependence
→ Tidal dissipation in stellar radiation zones and in planets
→ Differential rotation and magnetism




Thank
you !